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2009

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V. M. Abazov

Joint Institute for Nuclear Research, Dubna, Russia

Kenneth A. Bloom

University of Nebraska-Lincoln, kbloom2@unl.edu

Gregory Snow

University of Nebraska-Lincoln, gsnow1@unl.edu

D0 Collaboration

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Abazov, V. M.; Bloom, Kenneth A.; Snow, Gregory; and Collaboration, D0, "Search for Long-Lived Charged Massive Particles with the D0 Detector" (2009). *Kenneth Bloom Publications*. 309.

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Search for Long-Lived Charged Massive Particles with the D0 Detector

V. M. Abazov,³⁶ B. Abbott,⁷⁵ M. Abolins,⁶⁵ B. S. Acharya,²⁹ M. Adams,⁵¹ T. Adams,⁴⁹ E. Aguilo,⁶ M. Ahsan,⁵⁹ G. D. Alexeev,³⁶ G. Alkhazov,⁴⁰ A. Alton,^{64,*} G. Alverson,⁶³ G. A. Alves,² M. Anastasoia,³⁵ L. S. Ancu,³⁵ T. Andeen,⁵³ B. Andrieu,¹⁷ M. S. Anzels,⁵³ M. Aoki,⁵⁰ Y. Arnoud,¹⁴ M. Arov,⁶⁰ M. Arthaud,¹⁸ A. Askew,⁴⁹ B. Åsman,⁴¹ A. C. S. Assis Jesus,³ O. Atramentov,⁴⁹ C. Avila,⁸ F. Badaud,¹³ L. Bagby,⁵⁰ B. Baldin,⁵⁰ D. V. Bandurin,⁵⁹ P. Banerjee,²⁹ S. Banerjee,²⁹ E. Barberis,⁶³ A.-F. Barfuss,¹⁵ P. Bargassa,⁸⁰ P. Baringer,⁵⁸ J. Barreto,² J. F. Bartlett,⁵⁰ U. Bassler,¹⁸ D. Bauer,⁴³ S. Beale,⁶ A. Bean,⁵⁸ M. Begalli,³ M. Begel,⁷³ C. Belanger-Champagne,⁴¹ L. Bellantoni,⁵⁰ A. Bellavance,⁵⁰ J. A. Benitez,⁶⁵ S. B. Beri,²⁷ G. Bernardi,¹⁷ R. Bernhard,²³ I. Bertram,⁴² M. Besançon,¹⁸ R. Beuselinck,⁴³ V. A. Bezzubov,³⁹ P. C. Bhat,⁵⁰ V. Bhatnagar,²⁷ C. Biscarat,²⁰ G. Blazey,⁵² F. Blekman,⁴³ S. Blessing,⁴⁹ K. Bloom,⁶⁷ A. Boehnlein,⁵⁰ D. Boline,⁶² T. A. Bolton,⁵⁹ E. E. Boos,³⁸ G. Borissov,⁴² T. Bose,⁷⁷ A. Brandt,⁷⁸ R. Brock,⁶⁵ G. Brooijmans,⁷⁰ A. Bross,⁵⁰ D. Brown,⁸¹ X. B. Bu,⁷ N. J. Buchanan,⁴⁹ D. Buchholz,⁵³ M. Buehler,⁸¹ V. Buescher,²² V. Bunichev,³⁸ S. Burdin,^{42,†} T. H. Burnett,⁸² C. P. Buszello,⁴³ J. M. Butler,⁶² P. Calfayan,²⁵ S. Calvet,¹⁶ J. Cammin,⁷¹ M. A. Carrasco-Lizarraga,³³ E. Carrera,⁴⁹ W. Carvalho,³ B. C. K. Casey,⁵⁰ H. Castilla-Valdez,³³ S. Chakrabarti,¹⁸ D. Chakraborty,⁵² K. M. Chan,⁵⁵ A. Chandra,⁴⁸ E. Cheu,⁴⁵ F. Chevallier,¹⁴ D. K. Cho,⁶² S. Choi,³² B. Choudhary,²⁸ L. Christofek,⁷⁷ T. Christoudias,⁴³ S. Cihangir,⁵⁰ D. Claes,⁶⁷ J. Clutter,⁵⁸ M. Cooke,⁵⁰ W. E. Cooper,⁵⁰ M. Corcoran,⁸⁰ F. Couderc,¹⁸ M.-C. Cousinou,¹⁵ S. Crépe-Renaudin,¹⁴ V. Cuplov,⁵⁹ D. Cutts,⁷⁷ M. Ćwiok,³⁰ H. da Motta,² A. Das,⁴⁵ G. Davies,⁴³ K. De,⁷⁸ S. J. de Jong,³⁵ E. De La Cruz-Burelo,³³ C. De Oliveira Martins,³ K. DeVaughan,⁶⁷ F. Déliot,¹⁸ M. Demarteau,⁵⁰ R. Demina,⁷¹ D. Denisov,⁵⁰ S. P. Denisov,³⁹ S. Desai,⁵⁰ H. T. Diehl,⁵⁰ M. Diesburg,⁵⁰ A. Dominguez,⁶⁷ T. Dorland,⁸² A. Dubey,²⁸ L. V. Dudko,³⁸ L. Dufлот,¹⁶ S. R. Dugad,²⁹ D. Duggan,⁴⁹ A. Duperrin,¹⁵ J. Dyer,⁶⁵ A. Dyshkant,⁵² M. Eads,⁶⁷ D. Edmunds,⁶⁵ J. Ellison,⁴⁸ V. D. Elvira,⁵⁰ Y. Enari,⁷⁷ S. Eno,⁶¹ P. Ermolov,^{38,§§} H. Evans,⁵⁴ A. Evdokimov,⁷³ V. N. Evdokimov,³⁹ A. V. Ferapontov,⁵⁹ T. Ferbel,⁷¹ F. Fiedler,²⁴ F. Filthaut,³⁵ W. Fisher,⁵⁰ H. E. Fisk,⁵⁰ M. Fortner,⁵² H. Fox,⁴² S. Fu,⁵⁰ S. Fuess,⁵⁰ T. Gadfort,⁷⁰ C. F. Galea,³⁵ C. Garcia,⁷¹ A. Garcia-Bellido,⁷¹ V. Gavrilov,³⁷ P. Gay,¹³ W. Geist,¹⁹ W. Geng,^{15,65} C. E. Gerber,⁵¹ Y. Gershtein,^{49,‡} D. Gillberg,⁶ G. Ginter,⁷¹ B. Gómez,⁸ A. Goussiou,⁸² P. D. Grannis,⁷² H. Greenlee,⁵⁰ Z. D. Greenwood,⁶⁰ E. M. Gregores,⁴ G. Grenier,²⁰ Ph. Gris,¹³ J.-F. Grivaz,¹⁶ A. Grohsjean,²⁵ S. Grünendahl,⁵⁰ M. W. Grünwald,³⁰ F. Guo,⁷² J. Guo,⁷² G. Gutierrez,⁵⁰ P. Gutierrez,⁷⁵ A. Haas,⁷⁰ N. J. Hadley,⁶¹ P. Haefner,²⁵ S. Hagopian,⁴⁹ J. Haley,⁶⁸ I. Hall,⁶⁵ R. E. Hall,⁴⁷ L. Han,⁷ K. Harder,⁴⁴ A. Harel,⁷¹ J. M. Hauptman,⁵⁷ J. Hays,⁴³ T. Hebbeker,²¹ D. Hedin,⁵² J. G. Hegeman,³⁴ A. P. Heinson,⁴⁸ U. Heintz,⁶² C. Hensel,^{22,§} K. Herner,⁷² G. Hesketh,⁶³ M. D. Hildreth,⁵⁵ R. Hirosky,⁸¹ J. D. Hobbs,⁷² B. Hoeneisen,¹² M. Hohlfield,²² S. Hossain,⁷⁵ P. Houben,³⁴ Y. Hu,⁷² Z. Hubacek,¹⁰ V. Hynek,⁹ I. Iashvili,⁶⁹ R. Illingworth,⁵⁰ A. S. Ito,⁵⁰ S. Jabeen,⁶² M. Jaffré,¹⁶ S. Jain,⁷⁵ K. Jakobs,²³ C. Jarvis,⁶¹ R. Jesik,⁴³ K. Johns,⁴⁵ C. Johnson,⁷⁰ M. Johnson,⁵⁰ D. Johnston,⁶⁷ A. Jonckheere,⁵⁰ P. Jonsson,⁴³ A. Juste,⁵⁰ E. Kajfasz,¹⁵ D. Karmanov,³⁸ P. A. Kasper,⁵⁰ I. Katsanos,⁷⁰ D. Kau,⁴⁹ V. Kaushik,⁷⁸ R. Kehoe,⁷⁹ S. Kermiche,¹⁵ N. Khalatyan,⁵⁰ A. Khanov,⁷⁶ A. Kharchilava,⁶⁹ Y. M. Kharzheev,³⁶ D. Khatidze,⁷⁰ T. J. Kim,³¹ M. H. Kirby,⁵³ M. Kirsch,²¹ B. Klima,⁵⁰ J. M. Kohli,²⁷ J.-P. Konrath,²³ A. V. Kozelov,³⁹ J. Kraus,⁶⁵ T. Kuhl,²⁴ A. Kumar,⁶⁹ A. Kupco,¹¹ T. Kurča,²⁰ V. A. Kuzmin,³⁸ J. Kvita,⁹ F. Lacroix,¹³ D. Lam,⁵⁵ S. Lammers,⁷⁰ G. Landsberg,⁷⁷ P. Lebrun,²⁰ W. M. Lee,⁵⁰ A. Leflat,³⁸ J. Lellouch,¹⁷ J. Li,^{78,§§} L. Li,⁴⁸ Q. Z. Li,⁵⁰ S. M. Lietti,⁵ J. K. Lim,³¹ J. G. R. Lima,⁵² D. Lincoln,⁵⁰ J. Linnemann,⁶⁵ V. V. Lipaev,³⁹ R. Lipton,⁵⁰ Y. Liu,⁷ Z. Liu,⁶ A. Lobodenko,⁴⁰ M. Lokajicek,¹¹ P. Love,⁴² H. J. Lubatti,⁸² R. Luna-Garcia,^{33,||} A. L. Lyon,⁵⁰ A. K. A. Maciel,² D. Mackin,⁸⁰ R. J. Madaras,⁴⁶ P. Mättig,²⁶ C. Magass,²¹ A. Magerkurth,⁶⁴ P. K. Mal,⁸² H. B. Malbouisson,³ S. Malik,⁶⁷ V. L. Malyshev,³⁶ Y. Maravin,⁵⁹ B. Martin,¹⁴ R. McCarthy,⁷² M. M. Meijer,³⁵ A. Melnitchouk,⁶⁶ L. Mendoza,⁸ P. G. Mercadante,⁵ M. Merkin,³⁸ K. W. Merritt,⁵⁰ A. Meyer,²¹ J. Meyer,^{22,§} J. Mitrevski,⁷⁰ R. K. Mommsen,⁴⁴ N. K. Mondal,²⁹ R. W. Moore,⁶ T. Moulík,⁵⁸ G. S. Muanza,¹⁵ M. Mulhearn,⁷⁰ O. Mundal,²² L. Mundim,³ E. Nagy,¹⁵ M. Naimuddin,⁵⁰ M. Narain,⁷⁷ N. A. Naumann,³⁵ H. A. Neal,⁶⁴ J. P. Negret,⁸ P. Neustroev,⁴⁰ H. Nilsen,²³ H. Nogima,³ S. F. Novaes,⁵ T. Nunnemann,²⁵ V. O'Dell,⁵⁰ D. C. O'Neil,⁶ G. Obrant,⁴⁰ C. Ochando,¹⁶ D. Onoprienko,⁵⁹ N. Oshima,⁵⁰ N. Osman,⁴³ J. Osta,⁵⁵ R. Otec,¹⁰ G. J. Otero y Garzón,⁵⁰ M. Owen,⁴⁴ P. Padley,⁸⁰ M. Pangilinan,⁷⁷ N. Parashar,⁵⁶ S.-J. Park,^{22,§} S. K. Park,³¹ J. Parsons,⁷⁰ R. Partridge,⁷⁷ N. Parua,⁵⁴ A. Patwa,⁷³ G. Pawloski,⁸⁰ B. Penning,²³ M. Perfilov,³⁸ K. Peters,⁴⁴ Y. Peters,²⁶ P. Pétróff,¹⁶ M. Petteni,⁴³ R. Piegaia,¹ J. Piper,⁶⁵ M.-A. Pleier,²² P. L. M. Podesta-Lerma,^{33,||} V. M. Podstavkov,⁵⁰ Y. Pogorelov,⁵⁵ M.-E. Pol,² P. Polozov,³⁷ B. G. Pope,⁶⁵ A. V. Popov,³⁹ C. Potter,⁶ W. L. Prado da Silva,³ H. B. Prosper,⁴⁹ S. Protopopescu,⁷³ J. Qian,⁶⁴ A. Quadt,^{22,§} B. Quinn,⁶⁶ A. Rakitine,⁴² M. S. Rangel,² K. Ranjan,²⁸ P. N. Ratoff,⁴² P. Renkel,⁷⁹ P. Rich,⁴⁴ M. Rijssenbeek,⁷² I. Ripp-Baudot,¹⁹ F. Rizatdinova,⁷⁶ S. Robinson,⁴³ R. F. Rodrigues,³ M. Rominsky,⁷⁵ C. Royon,¹⁸ P. Rubinov,⁵⁰

R. Ruchti,⁵⁵ G. Safronov,³⁷ G. Sajot,¹⁴ A. Sánchez-Hernández,³³ M. P. Sanders,¹⁷ B. Sanghi,⁵⁰ G. Savage,⁵⁰ L. Sawyer,⁶⁰ T. Scanlon,⁴³ D. Schaile,²⁵ R. D. Schamberger,⁷² Y. Scheglov,⁴⁰ H. Schellman,⁵³ T. Schliephake,²⁶ S. Schlobohm,⁸² C. Schwanenberger,⁴⁴ A. Schwartzman,⁶⁸ R. Schwienhorst,⁶⁵ J. Sekaric,⁴⁹ H. Severini,⁷⁵ E. Shabalina,⁵¹ M. Shamim,⁵⁹ V. Shary,¹⁸ A. A. Shchukin,³⁹ R. K. Shivpuri,²⁸ V. Siccaldi,¹⁹ V. Simak,¹⁰ V. Sirotenko,⁵⁰ P. Skubic,⁷⁵ P. Slattery,⁷¹ D. Smirnov,⁵⁵ G. R. Snow,⁶⁷ J. Snow,⁷⁴ S. Snyder,⁷³ S. Söldner-Rembold,⁴⁴ L. Sonnenschein,¹⁷ A. Sopczak,⁴² M. Sosebee,⁷⁸ K. Soustruznik,⁹ B. Spurlock,⁷⁸ J. Stark,¹⁴ V. Stolin,³⁷ D. A. Stoyanova,³⁹ J. Strandberg,⁶⁴ S. Strandberg,⁴¹ M. A. Strang,⁶⁹ E. Strauss,⁷² M. Strauss,⁷⁵ R. Ströhmer,²⁵ D. Strom,⁵³ L. Stutte,⁵⁰ S. Sumowidagdo,⁴⁹ P. Svoisky,³⁵ A. Sznajder,³ A. Tanasijczuk,¹ W. Taylor,⁶ B. Tiller,²⁵ F. Tissandier,¹³ M. Titov,¹⁸ V. V. Tokmenin,³⁶ I. Torchiani,²³ D. Tsybychev,⁷² B. Tuchming,¹⁸ C. Tully,⁶⁸ P. M. Tuts,⁷⁰ R. Unalan,⁶⁵ L. Uvarov,⁴⁰ S. Uvarov,⁴⁰ S. Uzunyan,⁵² B. Vachon,⁶ P. J. van den Berg,³⁴ R. Van Kooten,⁵⁴ W. M. van Leeuwen,³⁴ N. Varelas,⁵¹ E. W. Varnes,⁴⁵ I. A. Vasilyev,³⁹ P. Verdier,²⁰ L. S. Vertogradov,³⁶ M. Verzocchi,⁵⁰ D. Vilanova,¹⁸ F. Villeneuve-Seguié,⁴³ P. Vint,⁴³ P. Vokac,¹⁰ M. Voutilainen,^{67,**} R. Wagner,⁶⁸ H. D. Wahl,⁴⁹ M. H. L. S. Wang,⁵⁰ J. Warchol,⁵⁵ G. Watts,⁸² M. Wayne,⁵⁵ G. Weber,²⁴ M. Weber,^{50,††} L. Welty-Rieger,⁵⁴ A. Wenger,^{23,‡‡} N. Wermes,²² M. Wetstein,⁶¹ A. White,⁷⁸ D. Wicke,²⁶ M. Williams,⁴² G. W. Wilson,⁵⁸ S. J. Wimpenny,⁴⁸ M. Wobisch,⁶⁰ D. R. Wood,⁶³ T. R. Wyatt,⁴⁴ Y. Xie,⁷⁷ C. Xu,⁶⁴ S. Yacoob,⁵³ R. Yamada,⁵⁰ W.-C. Yang,⁴⁴ T. Yasuda,⁵⁰ Y. A. Yatsunenko,³⁶ H. Yin,⁷ K. Yip,⁷³ H. D. Yoo,⁷⁷ S. W. Youn,⁵³ J. Yu,⁷⁸ C. Zeitnitz,²⁶ S. Zelitch,⁸¹ T. Zhao,⁸² B. Zhou,⁶⁴ J. Zhu,⁷² M. Zielinski,⁷¹ D. Zieminska,⁵⁴ A. Zieminski,^{54,§§} L. Zivkovic,⁷⁰ V. Zutshi,⁵² and E. G. Zverev³⁸

(D0 Collaboration)

¹Universidad de Buenos Aires, Buenos Aires, Argentina²LAFEX, Centro Brasileiro de Pesquisas Físicas, Rio de Janeiro, Brazil³Universidade do Estado do Rio de Janeiro, Rio de Janeiro, Brazil⁴Universidade Federal do ABC, Santo André, Brazil⁵Instituto de Física Teórica, Universidade Estadual Paulista, São Paulo, Brazil⁶University of Alberta, Edmonton, Alberta, Canada,

Simon Fraser University, Burnaby, British Columbia, Canada,

York University, Toronto, Ontario, Canada,

and McGill University, Montreal, Quebec, Canada

⁷University of Science and Technology of China, Hefei, People's Republic of China⁸Universidad de los Andes, Bogotá, Colombia⁹Center for Particle Physics, Charles University, Prague, Czech Republic¹⁰Czech Technical University, Prague, Czech Republic¹¹Center for Particle Physics, Institute of Physics, Academy of Sciences of the Czech Republic, Prague, Czech Republic¹²Universidad San Francisco de Quito, Quito, Ecuador¹³LPC, Université Blaise Pascal, CNRS/IN2P3, Clermont, France¹⁴LPSC, Université Joseph Fourier Grenoble 1, CNRS/IN2P3, Institut National Polytechnique de Grenoble, Grenoble, France¹⁵CPPM, Aix-Marseille Université, CNRS/IN2P3, Marseille, France¹⁶LAL, Université Paris-Sud, IN2P3/CNRS, Orsay, France¹⁷LPNHE, IN2P3/CNRS, Universités Paris VI and VII, Paris, France¹⁸CEA, Irfu, SPP, Saclay, France¹⁹IPHC, Université Louis Pasteur, CNRS/IN2P3, Strasbourg, France²⁰IPNL, Université Lyon 1, CNRS/IN2P3, Villeurbanne, France and Université de Lyon, Lyon, France²¹III. Physikalisches Institut A, RWTH Aachen University, Aachen, Germany²²Physikalisches Institut, Universität Bonn, Bonn, Germany²³Physikalisches Institut, Universität Freiburg, Freiburg, Germany²⁴Institut für Physik, Universität Mainz, Mainz, Germany²⁵Ludwig-Maximilians-Universität München, München, Germany²⁶Fachbereich Physik, University of Wuppertal, Wuppertal, Germany²⁷Panjab University, Chandigarh, India²⁸Delhi University, Delhi, India²⁹Tata Institute of Fundamental Research, Mumbai, India³⁰University College Dublin, Dublin, Ireland³¹Korea Detector Laboratory, Korea University, Seoul, Korea³²SungKyunKwan University, Suwon, Korea³³CINVESTAV, Mexico City, Mexico³⁴FOM-Institute NIKHEF and University of Amsterdam/NIKHEF, Amsterdam, The Netherlands

- ³⁵Radboud University Nijmegen/NIKHEF, Nijmegen, The Netherlands
³⁶Joint Institute for Nuclear Research, Dubna, Russia
³⁷Institute for Theoretical and Experimental Physics, Moscow, Russia
³⁸Moscow State University, Moscow, Russia
³⁹Institute for High Energy Physics, Protvino, Russia
⁴⁰Petersburg Nuclear Physics Institute, St. Petersburg, Russia
⁴¹Lund University, Lund, Sweden,
Royal Institute of Technology and Stockholm University, Stockholm, Sweden,
and Uppsala University, Uppsala, Sweden
⁴²Lancaster University, Lancaster, United Kingdom
⁴³Imperial College, London, United Kingdom
⁴⁴University of Manchester, Manchester, United Kingdom
⁴⁵University of Arizona, Tucson, Arizona 85721, USA
⁴⁶Lawrence Berkeley National Laboratory and University of California, Berkeley, California 94720, USA
⁴⁷California State University, Fresno, California 93740, USA
⁴⁸University of California, Riverside, California 92521, USA
⁴⁹Florida State University, Tallahassee, Florida 32306, USA
⁵⁰Fermi National Accelerator Laboratory, Batavia, Illinois 60510, USA
⁵¹University of Illinois at Chicago, Chicago, Illinois 60607, USA
⁵²Northern Illinois University, DeKalb, Illinois 60115, USA
⁵³Northwestern University, Evanston, Illinois 60208, USA
⁵⁴Indiana University, Bloomington, Indiana 47405, USA
⁵⁵University of Notre Dame, Notre Dame, Indiana 46556, USA
⁵⁶Purdue University Calumet, Hammond, Indiana 46323, USA
⁵⁷Iowa State University, Ames, Iowa 50011, USA
⁵⁸University of Kansas, Lawrence, Kansas 66045, USA
⁵⁹Kansas State University, Manhattan, Kansas 66506, USA
⁶⁰Louisiana Tech University, Ruston, Louisiana 71272, USA
⁶¹University of Maryland, College Park, Maryland 20742, USA
⁶²Boston University, Boston, Massachusetts 02215, USA
⁶³Northeastern University, Boston, Massachusetts 02115, USA
⁶⁴University of Michigan, Ann Arbor, Michigan 48109, USA
⁶⁵Michigan State University, East Lansing, Michigan 48824, USA
⁶⁶University of Mississippi, University, Mississippi 38677, USA
⁶⁷University of Nebraska, Lincoln, Nebraska 68588, USA
⁶⁸Princeton University, Princeton, New Jersey 08544, USA
⁶⁹State University of New York, Buffalo, New York 14260, USA
⁷⁰Columbia University, New York, New York 10027, USA
⁷¹University of Rochester, Rochester, New York 14627, USA
⁷²State University of New York, Stony Brook, New York 11794, USA
⁷³Brookhaven National Laboratory, Upton, New York 11973, USA
⁷⁴Langston University, Langston, Oklahoma 73050, USA
⁷⁵University of Oklahoma, Norman, Oklahoma 73019, USA
⁷⁶Oklahoma State University, Stillwater, Oklahoma 74078, USA
⁷⁷Brown University, Providence, Rhode Island 02912, USA
⁷⁸University of Texas, Arlington, Texas 76019, USA
⁷⁹Southern Methodist University, Dallas, Texas 75275, USA
⁸⁰Rice University, Houston, Texas 77005, USA
⁸¹University of Virginia, Charlottesville, Virginia 22901, USA
⁸²University of Washington, Seattle, Washington 98195, USA
(Received 26 September 2008; published 22 April 2009)

We search for long-lived charged massive particles using 1.1 fb^{-1} of data collected by the D0 detector at the Fermilab Tevatron $p\bar{p}$ Collider. Time-of-flight information is used to search for pair produced long-lived tau sleptons, gauginolike charginos, and Higgsino-like charginos. We find no evidence of a signal and set 95% C.L. cross section upper limits for staus, which vary from 0.31 to 0.04 pb for stau masses between 60 and 300 GeV. We also set lower mass limits of 206 GeV (171 GeV) for pair produced charged gauginos (Higgsinos).

Charged massive stable particles (CMSPs) are predicted by several extensions of the standard model. The term “stable” in this context refers to particles that live long enough to travel several meters long and escape a typical collider detector before decaying. The lightest tau slepton, or stau, is an example of such a particle, and is predicted in some gauge mediated supersymmetry breaking models [1]. If the stau decay is sufficiently suppressed, then the stau will be a CMSP candidate. The lightest chargino is another example of a CMSP. Anomaly mediated supersymmetry breaking models [2] or supersymmetric models that do not have gaugino mass unification can predict a long lifetime for the lightest chargino if its mass is within about 150 MeV of the lightest neutralino mass [3]. We explore two extreme cases, one where the chargino is mostly Higgsino and one where it is mostly gaugino.

Several collider experiments have performed searches for CMSPs. Studies at the CERN e^+e^- Collider (LEP) have resulted in lower mass limits of 97.5 GeV for stable sleptons [4], and 102.5 GeV for stable charginos [5]. A CDF Tevatron run I search set a cross section limit of $\mathcal{O}(1)$ pb for stable sleptons [6]. Complementary searches for neutral weakly interacting massive particles have also been performed by underground dark matter experiments [7].

In this Letter, we present a search for CMSPs produced directly in pairs. We do not consider CMSPs that result from cascade decays of heavier particles. The detector signature of pair produced CMSPs is rather striking. These weakly interacting particles are expected to traverse our entire detector, and should register in its outermost muon system. Additionally, owing to their large mass, these particles will travel substantially slower than beam produced muons, which travel near the speed of light.

Data used in this analysis were collected with the D0 detector [8] at the Fermilab Tevatron $p\bar{p}$ Collider at $\sqrt{s} = 1.96$ TeV between 2002 and 2006. They correspond to 1.1 fb^{-1} of integrated luminosity.

The D0 detector is a multipurpose detector well suited to a wide range of searches for new phenomena. The main components of the detector are an inner tracker, a liquid argon and uranium calorimeter, and a muon system. The inner tracker consists of a silicon microstrip detector (SMT) close to the beam line surrounded by a scintillating fiber detector. The muon system [9] resides beyond the calorimetry and consists of a layer of tracking detectors and scintillation trigger counters in front of 1.8 T iron toroids, followed by two similar layers after the toroids. Muon reconstruction at pseudorapidities [10] $|\eta| < 1$ relies on 10 cm wide drift tubes, while 1 cm mini drift tubes are used at $1 < |\eta| < 2$. Each scintillation counter registers a passing muon’s time, which can be used to calculate its speed [11]. The counters have a 2–4 ns time resolution, and their calibration is maintained run to run within 1 ns, relative to the event time determined from the accelerator clock.

The D0 detector uses a three-level trigger system to select data for offline analysis. CMSPs would appear as muons to the trigger system, so di-muon triggers were used to collect data for this analysis; and we use the term “muon” to refer to both real muons and CMSPs. Indeed, CMSPs are not distinguished from muons throughout the standard data collection and reconstruction, unless by virtue of their slow speed they arrive outside a muon trigger timing gate. The efficiency of the trigger gates is included in the calculated signal acceptance.

Muon candidates are reconstructed by finding tracks pointing to hit patterns in the muon system. We select events with exactly two muons, each of which satisfies quality criteria based on scintillator and drift tube information from the muon system and matches a track in the inner tracker. The muon candidates are also required to have transverse momenta, p_T , greater than 20 GeV, as measured with the central tracker. Events with muons from meson decays and other nonisolated muons are rejected by applying the following isolation criteria. At least one muon must have the sum of the p_T of all other tracks in a cone of radius $\mathcal{R} = \sqrt{(\Delta\phi)^2 + (\Delta\eta)^2} < 0.5$ around the muon direction less than 2.5 GeV. A similar isolation condition is applied for the total transverse energy measured in the calorimeter cells in a hollow cone of radius $0.1 < \mathcal{R} < 0.4$ around the muon direction; this energy must be less than 2.5 GeV.

A cosmic ray muon that passes through the detector can be reconstructed as two collinear muons. To reject these events, we require that the two muons must satisfy the pseudoacolinearity requirement $\Delta\alpha_{\mu\mu} = |\Delta\phi_{\mu\mu} + \Delta\theta_{\mu\mu} - 2\pi| > 0.05$. Moreover, since cosmic rays can arrive at times not correlated with the beam crossing, they can be misidentified as slow-moving particles. We also employ timing cuts which distinguish between outward going muons and inward traveling cosmic rays.

Two additional criteria are applied to reduce the background from muon candidates that do not originate from the primary vertex, such as those from cosmic rays, b decays, and beam halo. The distance of closest approach to the beam line (DCA), as measured in the transverse plane, for the track matched to the muon must be less than 0.02 cm for tracks with hits in the SMT and less than 0.2 cm for tracks without SMT hits. Finally, the difference in the z coordinates of the two muons at their DCAs is required to be less than 3 cm.

We determine the total signal acceptance using a combination of information from Monte Carlo (MC) simulation and the data. Signal samples for CMSP masses ranging between 60 and 300 GeV were generated with PYTHIA [12] using CTEQ6.1L [13] parton distribution functions (PDF), and processed with a GEANT-based [14] simulation of the D0 detector. These samples were reconstructed with the same software as the data. The specific model used for the stable stau is model line D in Ref. [15]. For long-lived

charginos, the specific models follow those described in Ref. [16]. In the present analysis only direct pair production is considered so that the exact values of the model parameters of the entire supersymmetric particle mass spectrum, relevant for cascade decays, are not important.

For each scintillator layer in which the reconstructed muon has a hit, the speed of this muon is calculated and expressed in units of the speed of light. The average speed \bar{v} is then obtained by taking the weighted average of the individual layer speeds. To ensure that the registered times in the muon detector are consistent, we compute a speed χ^2 from the individual layer speeds and their uncertainties. We require this $\chi^2/\text{d.o.f.}$ to be less than 4.7, a value derived from $Z \rightarrow \mu^+ \mu^-$ data. The transverse momenta of the pair produced CMSPs are expected to be approximately equal, and higher than those of beam produced muons. To reject tracks with poorly measured momentum, we require that the absolute value of the difference over the sum of the p_T of the two muon candidates in the event be less than 0.68, a value that is also derived using $Z \rightarrow \mu^+ \mu^-$ data.

Speed significance, defined as $(1 - \bar{v})/\sigma_{\bar{v}}$, is used to distinguish slow-moving particles from near light-speed muons. Here $\sigma_{\bar{v}}$ is the uncertainty in the average speed \bar{v} . We require that both candidate particles in the event should have positive speed significance.

In addition to time of flight, we use the invariant mass formed from the pair of muon candidates to separate signal

events from background. We calculate the invariant mass assuming the mass of each particle is that of a muon.

The only standard model background for this search comes from events which have, due to imperfect detector performance, anomalously large time of flight or mismeasured p_T that satisfy the selection criteria. Each of these measurements is independent of the other, since the p_T of the particle is measured in the central tracking system and the time of flight is recorded in the muon detector. Consequently, background events can be simulated by combining separate distributions of the invariant mass and of the speed significance product (the product of the values of speed significance of the two muon candidates). Events which pass all the selection criteria and have an invariant mass within the Z mass peak region (between 70 and 110 GeV) are used to model the speed significance product distribution for the background. The invariant mass distribution for the background is estimated from data events that have muon candidates with negative speed significance but pass all the other selection criteria. There is no overlap of events between these two data sets. Background events are then simulated by choosing a random value from each of the above two distributions, the invariant mass and the speed significance product, and are normalized to the number of Z and Drell-Yan events passing the selection criteria. The background and MC signal samples have very different distributions, as indicated in Fig. 1. They are combined using a joint likelihood to discriminate between background and signal. The likelihood discriminant cut values are chosen for each point by minimizing the expected 95% C.L. upper limit on the cross section calculated with a Bayesian limit method assuming a flat prior [17].

By construction, there is a correlation between data and the data-based simulated background in the speed significance distribution, but not in invariant mass. However, by randomly selecting from both templates, drawn from non-overlapping data sets, the variables are decorrelated for background while signal would show a correlation and peak more strongly at high likelihoods.

The signal acceptance, expected number of signal events, predicted number of background events, and the number of observed events are summarized in Table I for staus and charginos. The three models studied have different signal acceptances, reflecting the different CMSP kinematics of each model. The number of the observed events is consistent with the predicted background. A 95% C.L. upper limit on the pair production cross section is set for each mass point for the three models.

The systematic uncertainties in the background estimation arise mainly from the choice of the data events, whose invariant mass and speed significance product distributions are used to simulate the background. We varied the criteria used to select the data events, and the resulting difference in the predicted number of background events is taken as

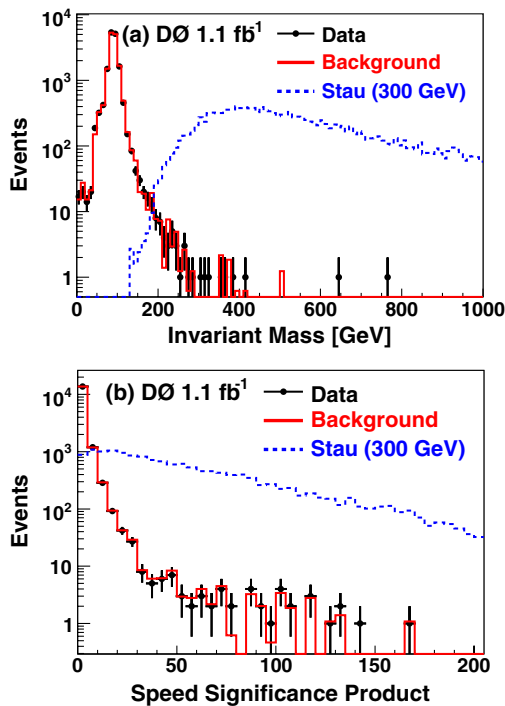


FIG. 1 (color online). Distributions of (a) the invariant mass and (b) speed significance product, for the simulated background (solid line), the 300 GeV stau signal (dotted line), and the data (as dots) passing the selection criteria.

TABLE I. Signal acceptance, expected number of signal events, predicted number of background events, and number of observed events for (a) staus, (b) gauginolike charginos, and (c) Higgsino-like charginos searches, as a function of the CMSP mass. The first uncertainty is statistical and the second is systematic.

Mass (GeV)	Signal acceptance ($\times 10^{-3}$)	Expected signal events	Predicted background	Observed events
(a) Stau				
60	$64 \pm 1 \pm 5$	4.7	$30.9 \pm 2.2 \pm 1.9$	38
80	$38 \pm 1 \pm 5$	1.1	$2.6 \pm 0.6 \pm 0.4$	1
100	$56 \pm 1 \pm 4$	0.7	$1.6 \pm 0.5 \pm 0.3$	1
150	$123 \pm 2 \pm 13$	0.3	$1.7 \pm 0.5 \pm 0.2$	1
200	$139 \pm 2 \pm 11$	0.1	$1.7 \pm 0.5 \pm 0.5$	1
250	$133 \pm 2 \pm 13$	0.01	$1.7 \pm 0.5 \pm 0.3$	1
300	$117 \pm 2 \pm 13$	0.004	$1.9 \pm 0.5 \pm 0.2$	2
(b) Gauginolike charginos				
60	$32 \pm 1 \pm 3$	445	$23.6 \pm 1.9 \pm 1.4$	24
80	$24 \pm 1 \pm 3$	85	$1.9 \pm 0.5 \pm 0.3$	1
100	$46 \pm 1 \pm 4$	65	$1.6 \pm 0.5 \pm 0.3$	1
150	$85 \pm 1 \pm 9$	20	$1.2 \pm 0.4 \pm 0.1$	1
200	$89 \pm 1 \pm 7$	5	$1.9 \pm 0.5 \pm 0.0$	1
250	$74 \pm 1 \pm 7$	1	$1.7 \pm 0.5 \pm 0.3$	1
300	$59 \pm 1 \pm 7$	0.2	$1.7 \pm 0.5 \pm 0.1$	2
(c) Higgsino-like charginos				
60	$29 \pm 1 \pm 2$	94	$17.9 \pm 1.7 \pm 1.1$	21
80	$24 \pm 1 \pm 3$	23	$1.6 \pm 0.5 \pm 0.3$	1
100	$49 \pm 1 \pm 4$	20	$1.6 \pm 0.5 \pm 0.3$	1
150	$89 \pm 1 \pm 9$	7	$1.4 \pm 0.5 \pm 0.1$	1
200	$96 \pm 1 \pm 8$	2	$1.9 \pm 0.5 \pm 0.0$	1
250	$81 \pm 1 \pm 8$	0.5	$1.7 \pm 0.5 \pm 0.3$	1
300	$64 \pm 1 \pm 7$	0.1	$1.7 \pm 0.5 \pm 0.1$	1

the size of the systematic uncertainty. The main signal acceptance uncertainties are those in object identification efficiencies, trigger efficiencies, MC simulation normalizations, and uncertainties related to the choice of PDF.

The masses and couplings are computed by SOFTSUSY [18], and the next-to-leading order (NLO) cross section is calculated with PROSPINO2.0 [19]. The renormalization and factorization scale uncertainty and the PDF uncertainty are added in quadrature to obtain the total uncertainty on the signal cross section. The calculated expected and observed

limits, the NLO cross section, and the uncertainty on the cross section are shown in Fig. 2 for varying stau and chargino masses. Using the nominal (nominal -1σ) values of the NLO cross section, lower mass limits of 206 (204) GeV at 95% C.L. are set for gauginolike charginos. For Higgsino-like charginos the limits are 171 (169) GeV. Although the present sensitivity is insufficient to test the model of pair produced staus, the cross section limits can be applied to the pair production of any CMSP candidate with similar kinematics.

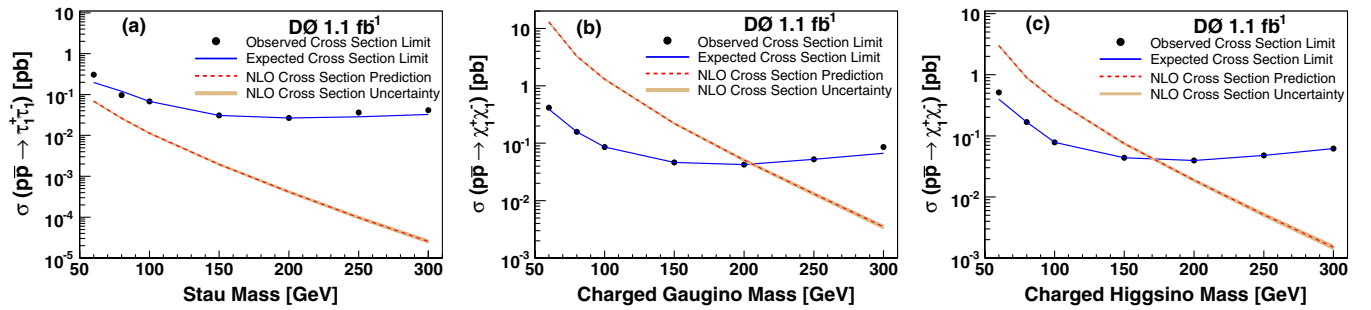


FIG. 2 (color online). The observed (dots) and expected (solid line) 95% cross section limits, the NLO production cross section (dashed line), and NLO cross section uncertainty (barely visible shaded band) as a function of (a) stau mass for stau pair production, (b) chargino mass for pair produced gauginolike charginos, and (c) chargino mass for pair produced Higgsino-like charginos.

In summary, we have performed a search for charged massive stable particles using 1.1 fb^{-1} of data collected by the D0 detector. We find no evidence of a signal and set 95% C.L. cross section limits on the pair production of stable staus and gauginolike and Higgsino-like charginos. The upper cross section limits vary from 0.31 to 0.04 pb for stau masses in the range 60–300 GeV. We use the nominal value of the theoretical cross section to set limits on the mass of pair produced charginos. We exclude stable gauginolike charginos with masses below 206 GeV and Higgsino-like charginos below 171 GeV. These are the most restrictive limits to date on the cross sections for CMSPs and the first published from the Tevatron Collider run II.

We thank the staffs at Fermilab and collaborating institutions, and acknowledge support from the DOE and NSF (U.S.A.); CEA and CNRS/IN2P3 (France); FASI, Rosatom, and RFBR (Russia); CNPq, FAPERJ, FAPESP, and FUNDUNESP (Brazil); DAE and DST (India); Colciencias (Colombia); CONACyT (Mexico); KRF and KOSEF (Korea); CONICET and UBACyT (Argentina); FOM (The Netherlands); STFC (United Kingdom); MSMT and GACR (Czech Republic); CRC Program, CFI, NSERC, and WestGrid Project (Canada); BMBF and DFG (Germany); SFI (Ireland); The Swedish Research Council (Sweden); CAS and CNSF (China); and the Alexander von Humboldt Foundation (Germany).

*Visitor from Augustana College, Sioux Falls, SD, USA.

†Visitor from The University of Liverpool, Liverpool, UK.

‡Visitor from Rutgers University, Piscataway, NJ, USA.

§Visitor from II. Physikalisches Institut, Georg-August-University, Göttingen, Germany.

||Visitor from Centro de Investigacion en Computacion-IPN, Mexico City, Mexico.

¶Visitor from ECFM, Universidad Autonoma de Sinaloa, Culiacán, Mexico.

**Visitor from Helsinki Institute of Physics, Helsinki, Finland.

††Visitor from Universität Bern, Bern, Switzerland.

‡‡Visitor from Universität Zürich, Zürich, Switzerland.

§§Deceased.

- [1] G. F. Giudice and R. Rattazzi, *Phys. Rep.* **322**, 419 (1999).
 [2] J. Gunion and S. Mrenna, *Phys. Rev. D* **62**, 015002 (2000).

- [3] C. Chen, M. Drees, and J. Gunion, *Phys. Rev. D* **55**, 330 (1997).
 [4] LEP SUSYWG, ALEPH, DELPHI, L3, and OPAL Collaborations, Note LEP SUSYWG/02-09.2, http://lepsusy.web.cern.ch/lepsusy/www/gmsb_summer02/lepgmsb.html.
 [5] LEP SUSYWG, ALEPH, DELPHI, L3, and OPAL Collaborations, Note LEP SUSYWG/02-05.1, http://lepsusy.web.cern.ch/lepsusy/www/stable_summer02/stable_208.html.
 [6] D. Acosta *et al.* (CDF Collaboration), *Phys. Rev. Lett.* **90**, 131801 (2003).
 [7] R. J. Gaitskell, *Annu. Rev. Nucl. Part. Sci.* **54**, 315 (2004).
 [8] V. Abazov *et al.* (D0 Collaboration), *Nucl. Instrum. Methods Phys. Res., Sect. A* **565**, 463 (2006).
 [9] V. Abazov *et al.*, *Nucl. Instrum. Methods Phys. Res., Sect. A* **552**, 372 (2005).
 [10] The D0 coordinate system is cylindrical with the z axis along the proton beam direction and the polar and azimuthal angles denoted as θ and ϕ , respectively. The pseudorapidity is defined as $\eta = -\ln[\tan(\theta/2)]$.
 [11] There are up to three layers of scintillation counters with each layer forming approximately the sides of a cube. The closest distance from the interaction point to each layer is 2.97, 4.45, and 5.7 m. The width of the timing gate ranges from 24 to 84 ns, depending on the layer.
 [12] T. Sjöstrand *et al.*, *Comput. Phys. Commun.* **135**, 238 (2001). We use PYTHIA 6.323.
 [13] J. Pumplin *et al.* (CTEQ Collaboration), *J. High Energy Phys.* 07 (2002) 012.
 [14] R. Brun *et al.*, CERN Program Library Long Writeup No. W5013, 1993.
 [15] S. P. Martin, S. Moretti, J. Qian, and G. W. Wilson, in *Proceedings of APS/DPF/DPB Summer Study on the Future of Particle Physics*, eConf C010630, P346 (2001). In this model, Λ is approximately equal to the stau mass, $M_m = 2\Lambda$, $N_5 = 3$, $\tan\beta = 15$, and $\mu > 0$; See also model SPS7 in B. C. Allanach *et al.*, *Eur. Phys. J. C* **25**, 113 (2002).
 [16] P. Abreu *et al.* (DELPHI Collaboration), *Eur. Phys. J. C* **11**, 1 (1999). In the present analysis, for gauginolike charginos, $\mu = 10^4$ GeV, M_2 is approximately equal to the mass of the lightest chargino, and $M_1 = 3M_2$. For Higgsino-like chargino production, $M_1 = M_2 = 10^5$ GeV and μ is approximately equal to the mass of the lightest chargino. In both models, $\tan\beta = 15$, $M_3 = 500$ GeV, and $M_{\tilde{g}} = 800$ GeV.
 [17] I. Bertram *et al.*, Fermilab Report No. Fermilab-TM-2104, 2000.
 [18] B. C. Allanach, *Comput. Phys. Commun.* **143**, 305 (2002).
 [19] W. Beenakker *et al.*, *Phys. Rev. Lett.* **83**, 3780 (1999).